LIGHT AXINO DARK MATTER FROM FREEZE-IN PRODUCTION

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based on
arXiv:1707.02077, 1707.06418
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OVERVIEW

- ***"Warm" Dark Matter: m~keV***

  - signal may be seen; 3.5 keV X-ray observation (7 keV DM decay).

  - Lyman-alpha forest constraints: pushing keV DM to heavier region.

- ***"Less" Warm Dark Matter***

  - Fixing DM mass 7 keV,

  - Non-standard DM production; "Colder" warm dark matter,

  - Entropy injection also helps.

- ***Light Axino DM***

  - keV scale DM

  - Freeze-in Production

  - Entropy from "Saxion"

AXINO

- ***Supersymmetry+Peccei-Quinn Symmetry:***

  - SUSY solves the gauge hierarchy & PQ solves strong CP

  - Dark Matter candidates: neutralino & axion

- ***Axino Dark Matter:***

  - fermionic SUSY partner of axion

  - becomes massive with SUSY breaking

  - keV mass is possible: (warm) *Dark Matter*

- ***Signal?***

  - too weak to be detected: suppressed by $\sim 10^9$ GeV

  - *what if it decays?*

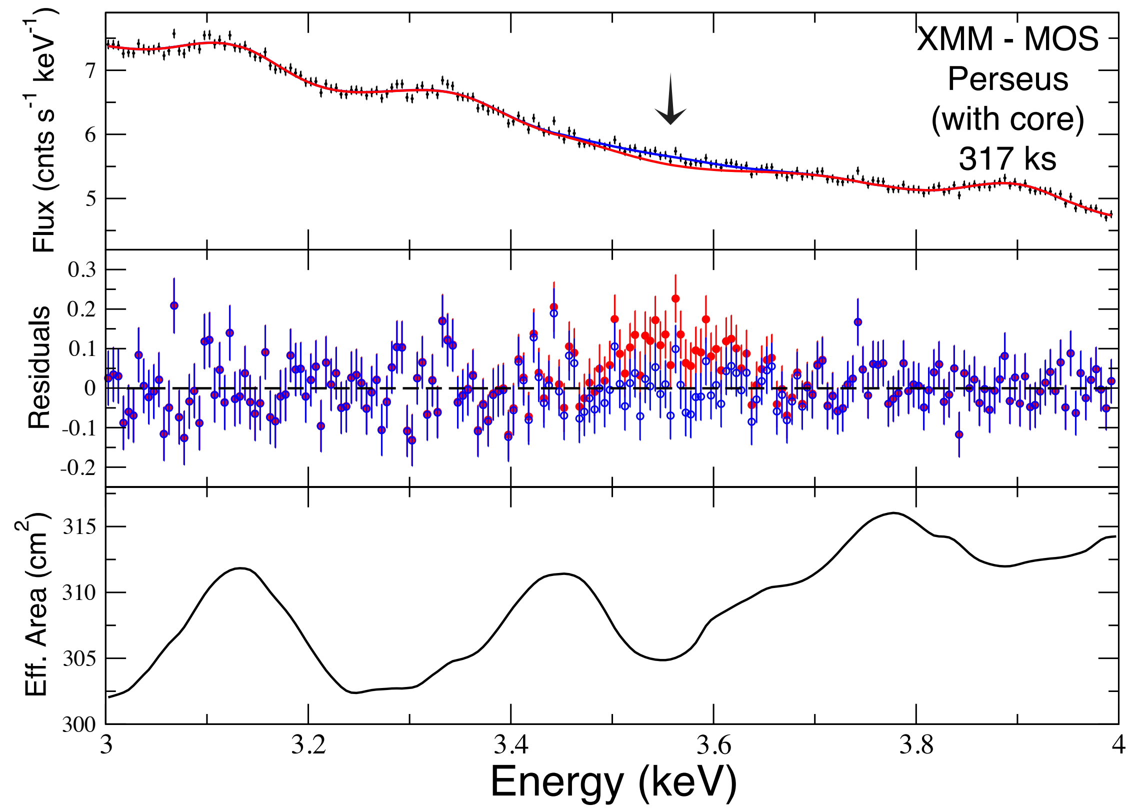
  - affects *matter power spectrum?*

X-RAY LINE?

- 3.5 keV X-ray line excess

  Bulbul, Markevitch, Foster, Smith, Lowenstein, Randall(2014)

  Boyarsky, Ruchayskiy, Iakubovskyi, Franse (2014)



- Criticism:

  K XVIII explanation

  No signal from dSph, stacked galaxies and groups, M31

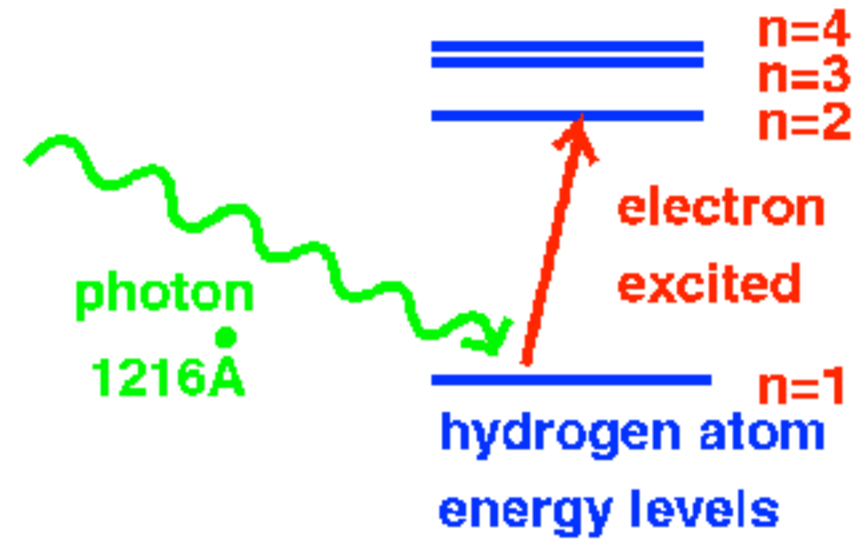
  Morphology study

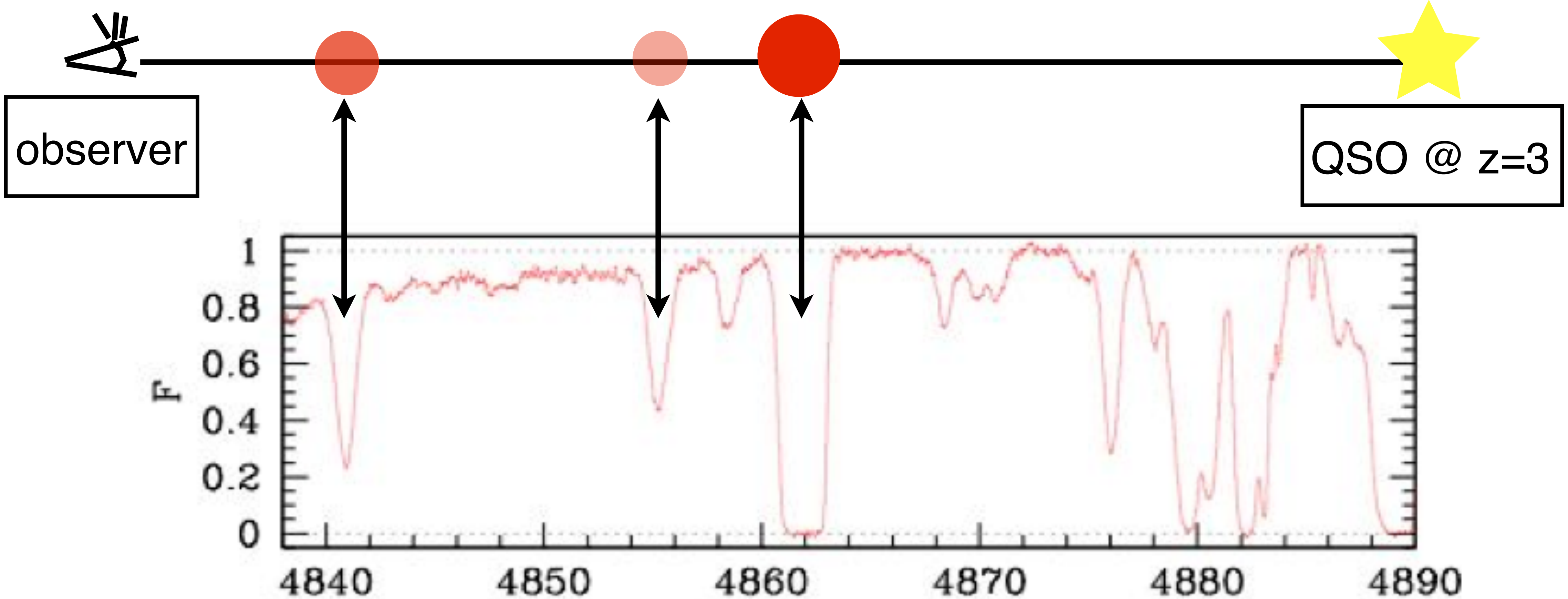
- In this study,

  We take 7 keV as a benchmark point (and a possible signal) for light axino study.

LYMAN-ALPHA FOREST

absorption intensity/frequency
↔ HI distribution along the line-of-sight





slide by A. Kamada

LYMAN-ALPHA FOREST

- Improving constraints on "warm dark matter mass"

$m_{\rm WDM} \gtrsim 2.0~{\rm keV}$ Viel, Lesgourgues, Haehnelt, Matarrese, Riotto (2005)

$m_{\rm WDM} \gtrsim 3.3~{\rm keV}$ Viel, Becker, Bolton, Haehnelt (2013)

$m_{\rm WDM} \gtrsim 4.09~{\rm keV}$ Baur, Palanque-Delabrouille, Yche, Magneville, Viel (2016)

$m_{\rm WDM} \gtrsim 5.3~{\rm keV}$ Iršič et al. (2017)

- ***Warm Dark Matter***

Ly-alpha constraints assume the Fermi-Dirac dist. and observed DM density

$$\Omega_{\rm WDM}h^2 = \left(\frac{m_{\rm WDM}}{94\,{\rm eV}}\right)\left(\frac{T_{\rm WDM}}{T_\nu}\right)^3 = 7.5\left(\frac{m_{\rm WDM}}{7\,{\rm keV}}\right)\left(\frac{106.75}{g_*^{\rm WDM}}\right)$$

$g_*^{\rm WDM} \sim 7000$ or $\Delta \sim 70$?

➡ need *linear matter power spectrum*

LYMAN-ALPHA FOREST

- Improving constraints on "warm dark matter mass"

$m_{\rm WDM} \gtrsim 2.0$ keV Viel Lesgourgues Haehnelt Matarrese Riotto (2005)

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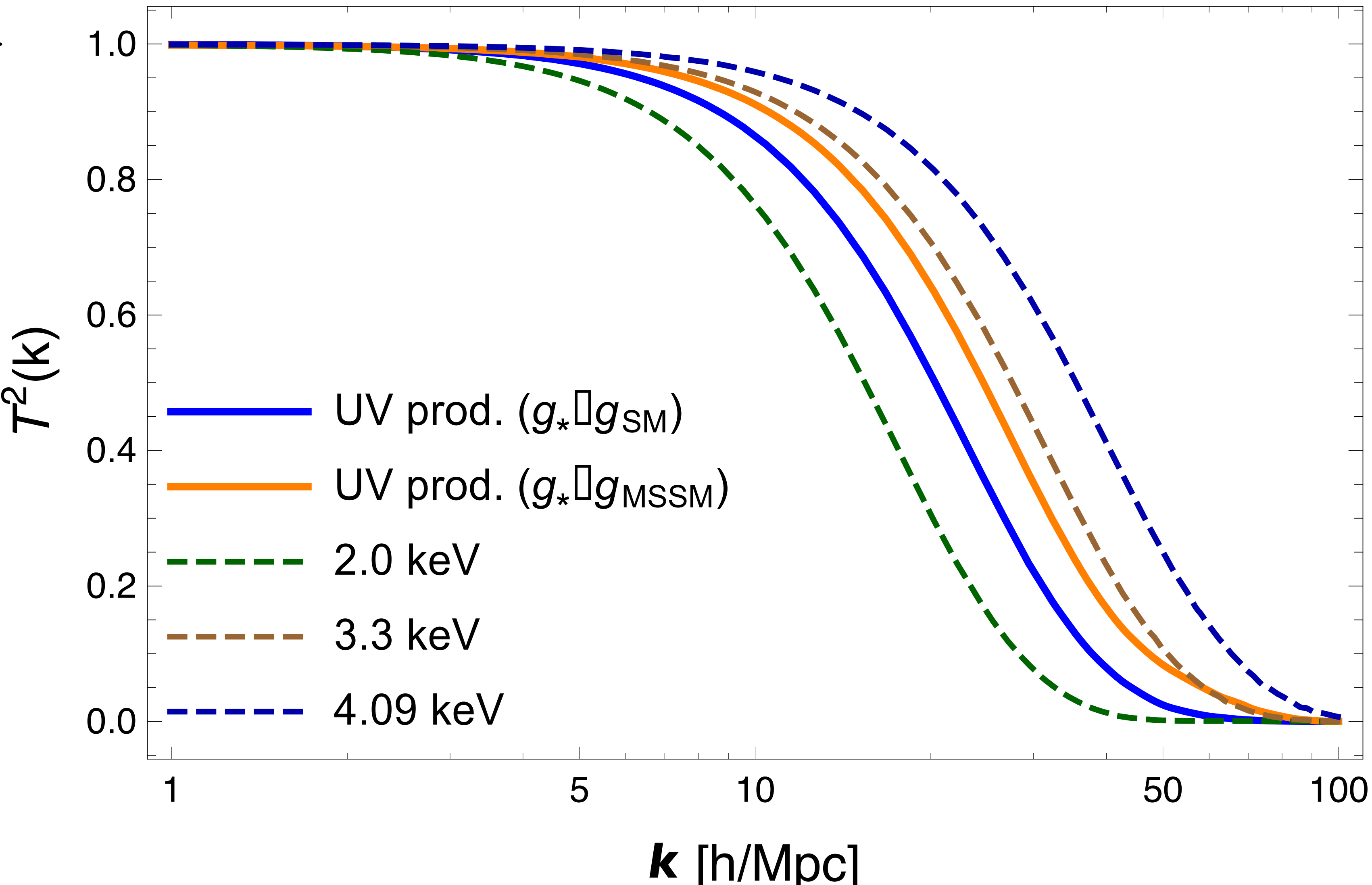

- ***Warm***

Ly-alpha ... observed

DM den

$\Omega_{\rm W}$ ...)

$g_*^{\rm WDM} \sim 70$

➡ need *linear matter power spectrum*

OUTLINE

1. Introduction
2. SUSY DFSZ model
3. Freeze-in Axinos
4. Linear Matter Power Spectrum
5. Conclusions

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AXINO IN SUSY

- ***SUSY+PQ***

motivated by both gauge hierarchy and strong CP

Goldstone axion is supersymmetrized

$$a \quad \longrightarrow \quad A = \frac{1}{\sqrt{2}}(s+ia) + \sqrt{2}\theta\tilde{a} + \theta^2 F_A$$

- ***Properties***

massless ➡ massive by SUSY breaking $m_{\tilde{a}} \sim m_{3/2}$

But in some models, $m_{\tilde{a}} \sim \mathcal{O}(\mathrm{keV})$

Tamvakis, Wyler; Nieves; Goto, Yamaguchi; Chun, Kim, Nilles; Chun, Lukas

couplings: suppressed by PQ scale $\gtrsim 10^9 \ \mathrm{GeV}$

➡ *Feebly Interacting Massive Particle*

or SuperWIMP

MODEL

- ***SUSY DFSZ model***

$$W_{\rm DFSZ} = \frac{y_0}{M_*} X^2 H_u H_d \,,$$

~~PQ~~

$$X = \frac{v_{\rm PQ}}{\sqrt{2}} e^{A/v_{\rm PQ}} \,,$$

- generates **mu-term** $\mu \sim \dfrac{y_0 v_{\rm PQ}^2}{2M_*}$

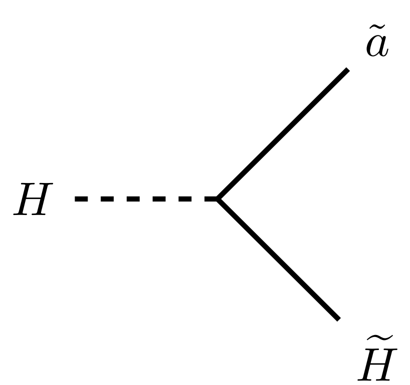
Kim, Nilles (1984)

"effectively dimensionless"

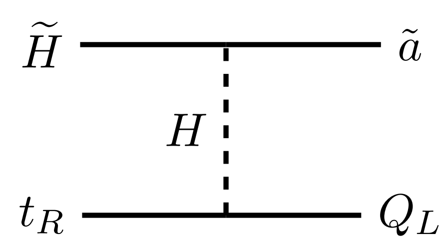
- **axino interaction** $\dfrac{2\mu}{v_{\rm PQ}} A H_u H_d$

$$\frac{\mu}{v_{\rm PQ}} \sim 10^{-8}$$

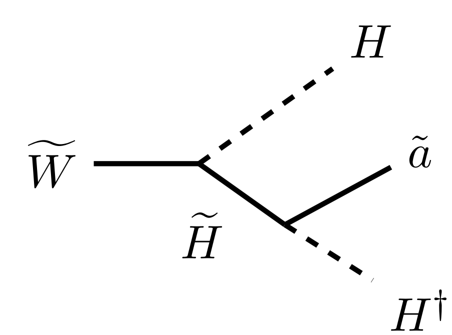
"***Freeze-in***" production Hall, Jedamzik, March-Russell, West (2009)



2-body decay



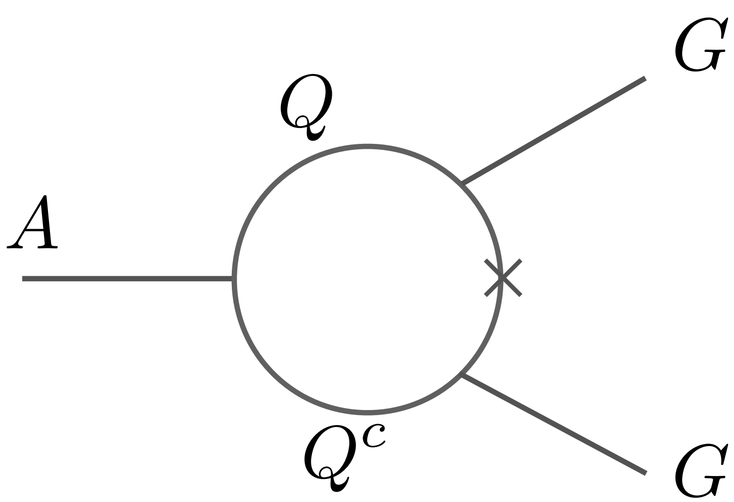
2-to-2



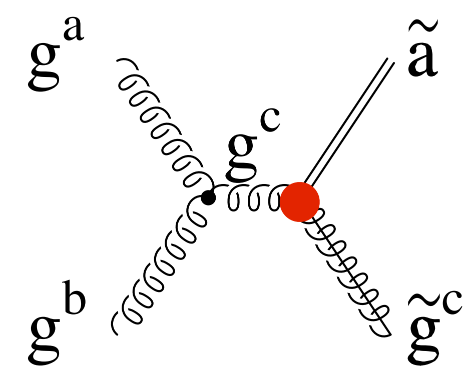
3-body decay

UV PRODUCTION?

- ElectroWeak Symmetry Breaking & SM quark loops



$$\sim \frac{g_s^2}{32\pi^2}\frac{1}{F_a}\bar{\tilde{a}}\gamma_5\sigma^{\mu\nu}\tilde{g}^a\partial_\mu G^a_\nu$$



- 1PI amplitude: suppressed $\mathcal{A}_{\rm 1PI} \propto \frac{m_Q^2}{E^2}$ KJB, Choi, Im (2011)

- Above the weak scale (or before EW phase transition), ***UV production is negligible.***

R-PARITY VIOLATION

- PQ charge $Q_{\rm PQ}\{X, H_u, H_d, L_i\} = \{-1, 1, 1, 2\}$

Chun (1999); Choi, Chun, Hwang (2001); Chun, Kim (2006)

$$W_{\rm bRPV} = \frac{y_i'}{M_*^2} X^3 L_i H_u \simeq \mu_i' \left(1 + \frac{3A}{v_{\rm PQ}}\right) L_i H_u,$$

- Bilinear RPV induces ***axino-neutrino mixing***

$$|\theta| \simeq \frac{\mu' v_u}{m_{\tilde a} v_{\rm PQ}} \simeq 10^{-5} \left(\frac{\mu'}{4\,{\rm MeV}}\right) \left(\frac{7\,{\rm keV}}{m_{\tilde a}}\right) \left(\frac{10^{10}\,{\rm GeV}}{v_{\rm PQ}}\right)$$

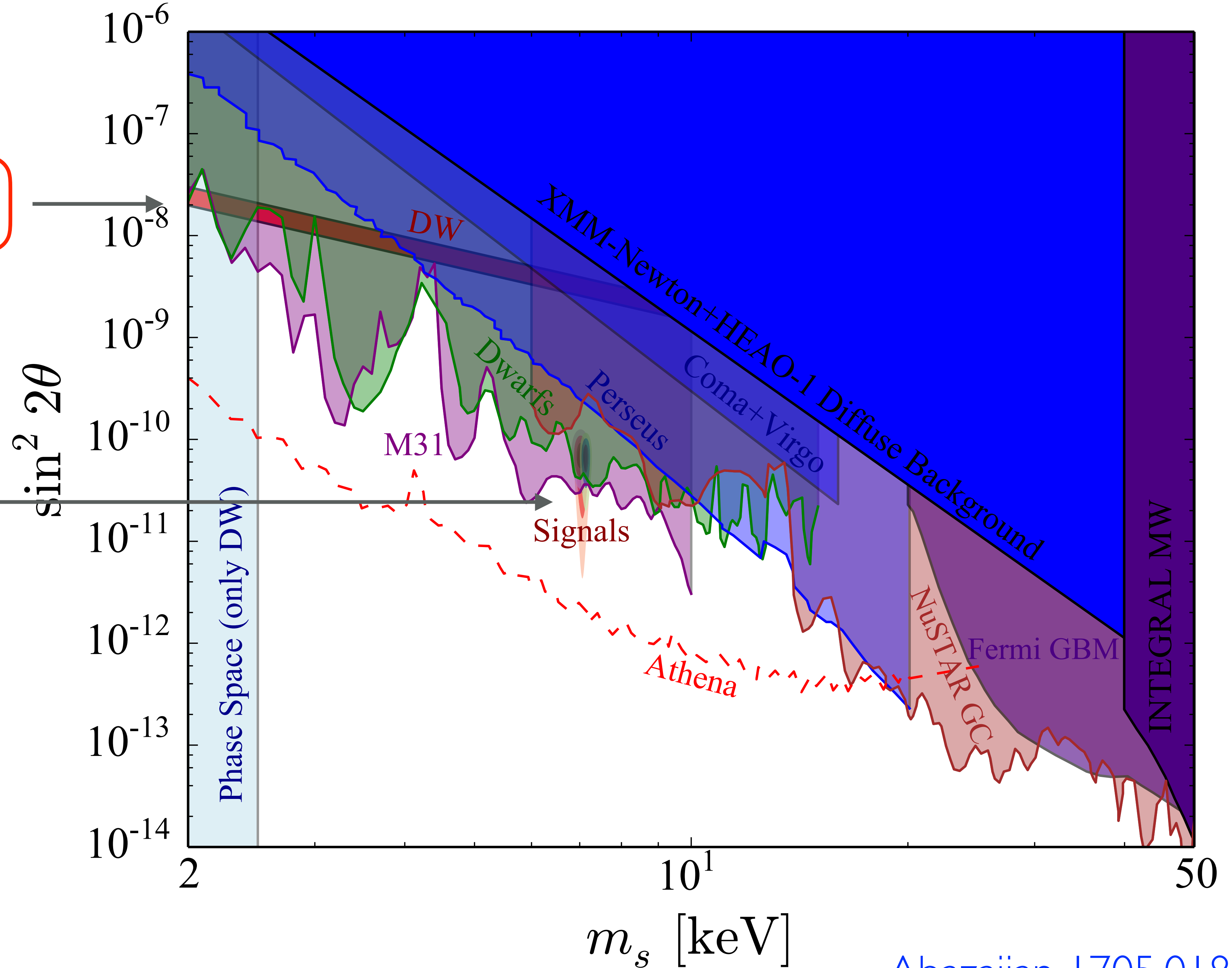
- ***Axino as a sterile neutrino***

  - decays into neutrino & *photon*
  - produced via Dodelson-Widrow mechanism

STERILE NEUTRINO

Dodelson-Widrow

X-ray line



Abazajian 1705.01837

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BOLTZMANN EQ.

- Boltzmann eq. for phase space distribution

$$\frac{df_{\tilde{a}}(t,p)}{dt} = \frac{\partial f_{\tilde{a}}(t,p)}{\partial t} - \frac{\dot{a}(t)}{a(t)} p \frac{\partial f_{\tilde{a}}(t,p)}{\partial p} = \frac{1}{E_{\tilde{a}}} C(t,p)$$

$f_{\tilde{a}}(t,p)$: phase space distribution

$a(t)$: cosmic scale factor

$E_{\tilde{a}}$: axino energy $\quad C(t,p)$: collision term

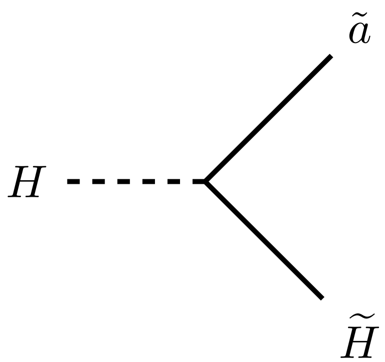
- Negligible $f_{\tilde{a}}(t,p)$ in the collision term

$$f_{\tilde{a}}(t_f,p) = \int_{t_i}^{t_f} dt \frac{1}{E_{\tilde{a}}} C\left(t, \frac{a(t_f)}{a(t)} p\right)$$

collision terms determine distribution

COLLISION TERMS

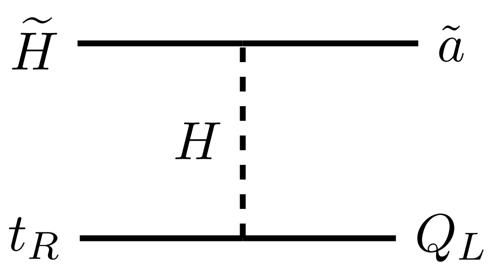
- 2-body decay



$$p_{\tilde{a}} \sim m_H/2 \qquad \text{vs.} \qquad T \sim m_H$$

or Higgsino decay

- 2-to-2 scattering

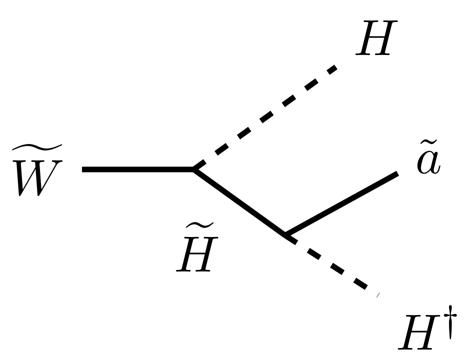


+ s-channel

$$p_{\tilde{a}} \sim \sqrt{s} \qquad \text{vs.} \qquad T \sim \sqrt{s}$$

different dists. for s- and t-channel

- 3-body decay

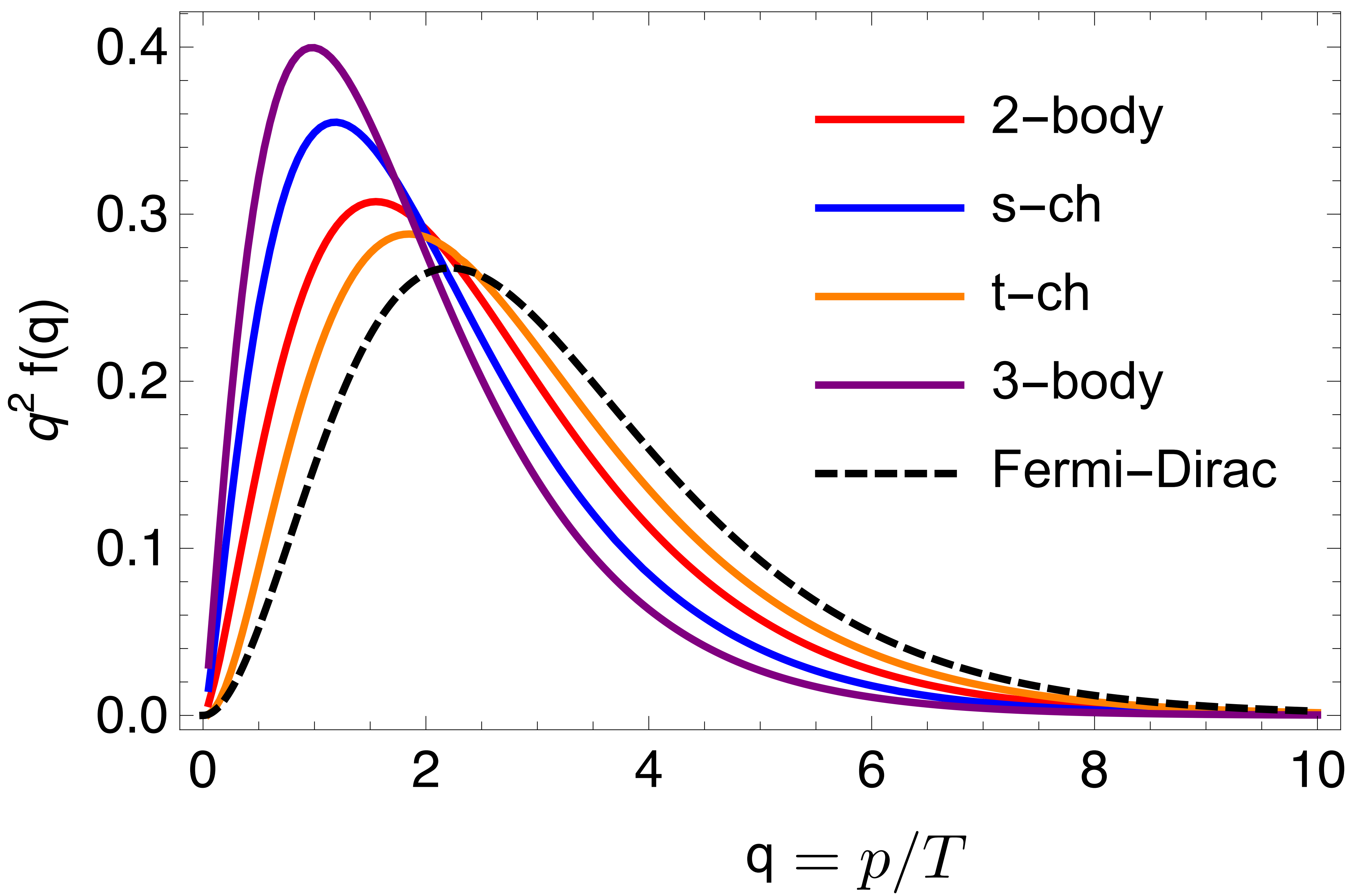


$$p_{\tilde{a}} \sim m_{\widetilde{W}}/3 \qquad \text{vs.} \qquad T \sim m_{\widetilde{W}}$$

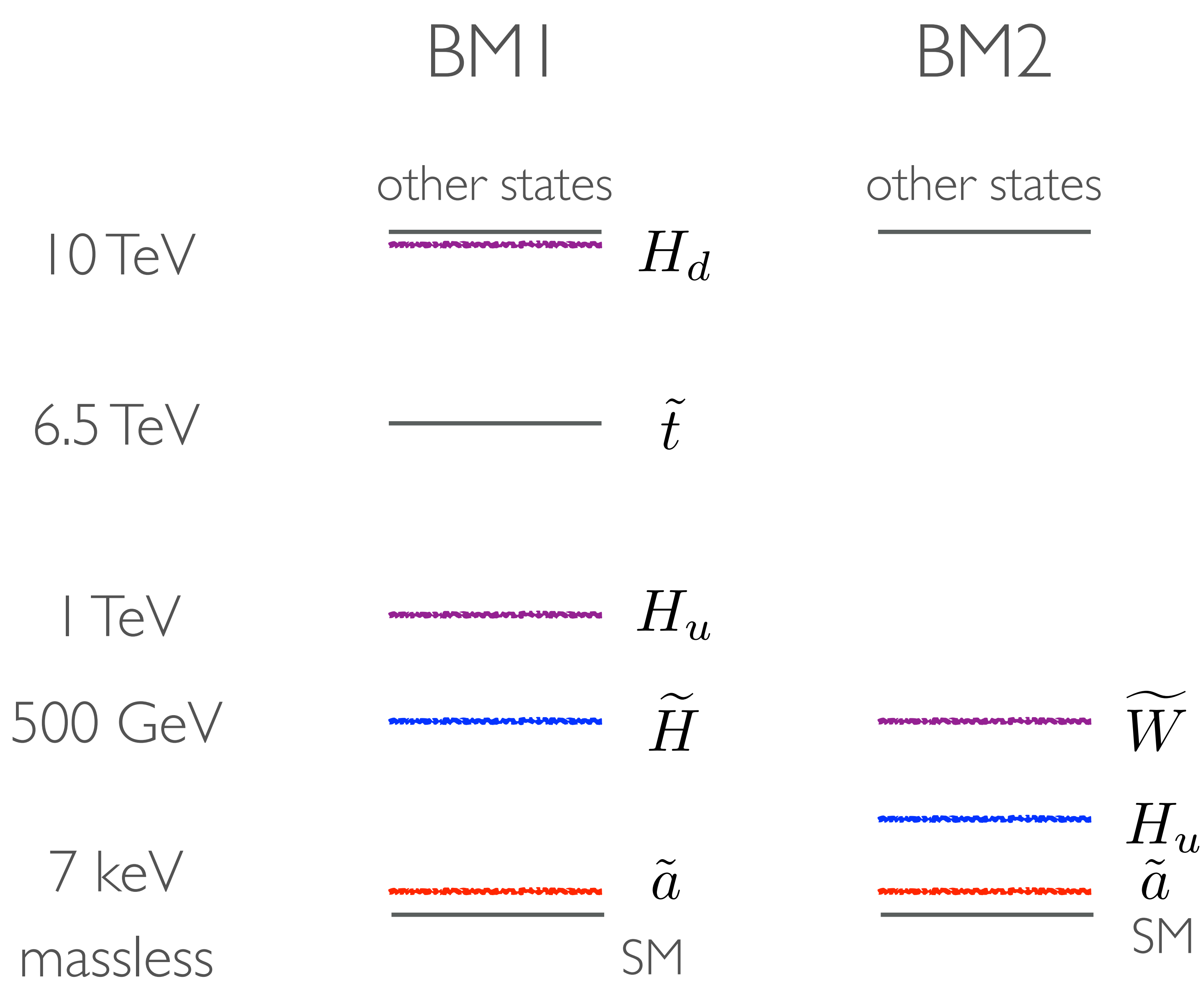
comparable to scattering cross section

e.g. $\widetilde{W} + H \to H + \tilde{a}$

DISTRIBUTION



BENCHMARK SCENARIOS



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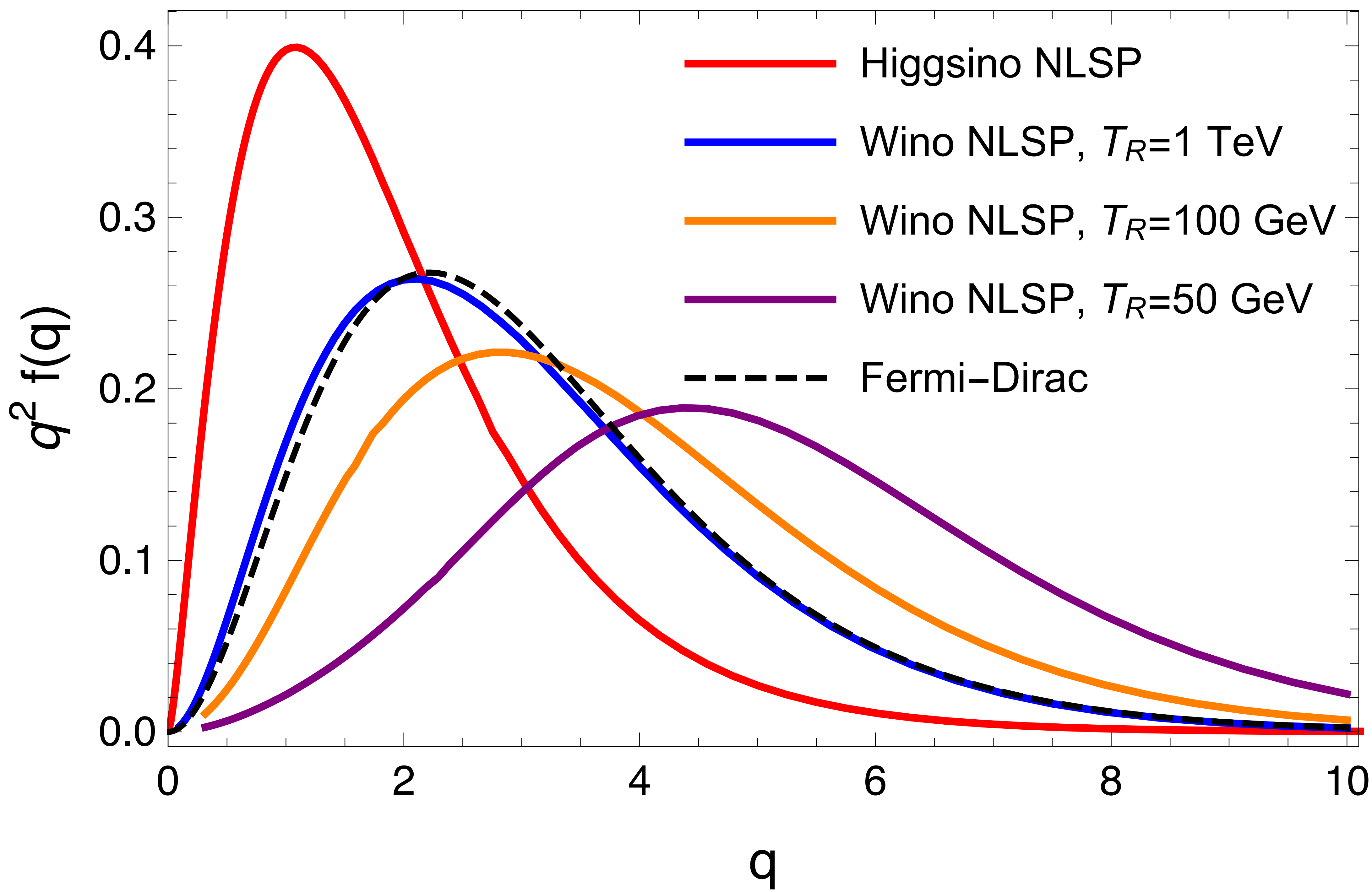

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SQUARED TRANSFER FUNC.

- Ly-alpha bounds: assuming FD dist. and observed DM density

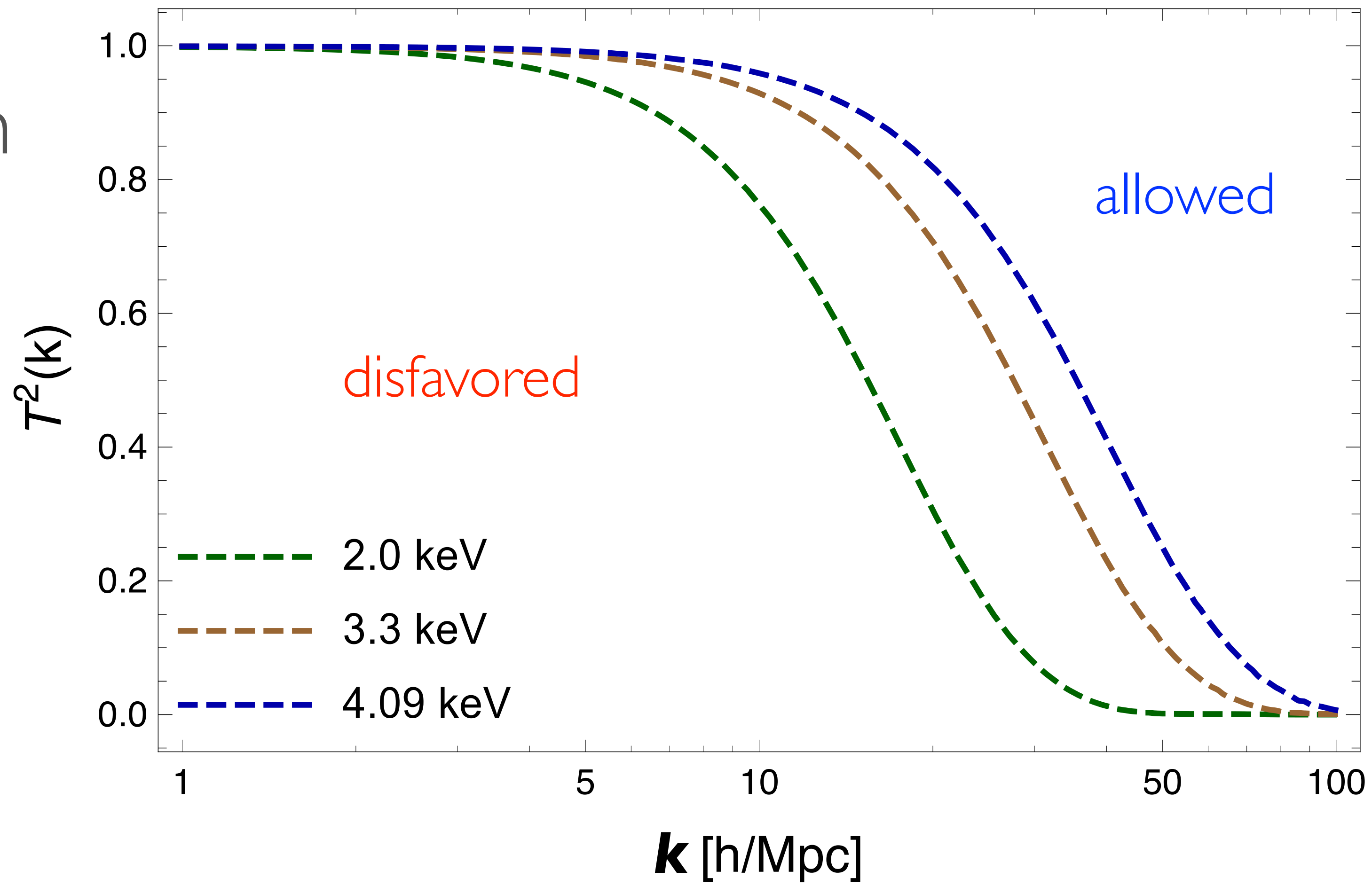
  ➡ constrain warm DM mass

- Freeze-in Axino: non-FD dist.

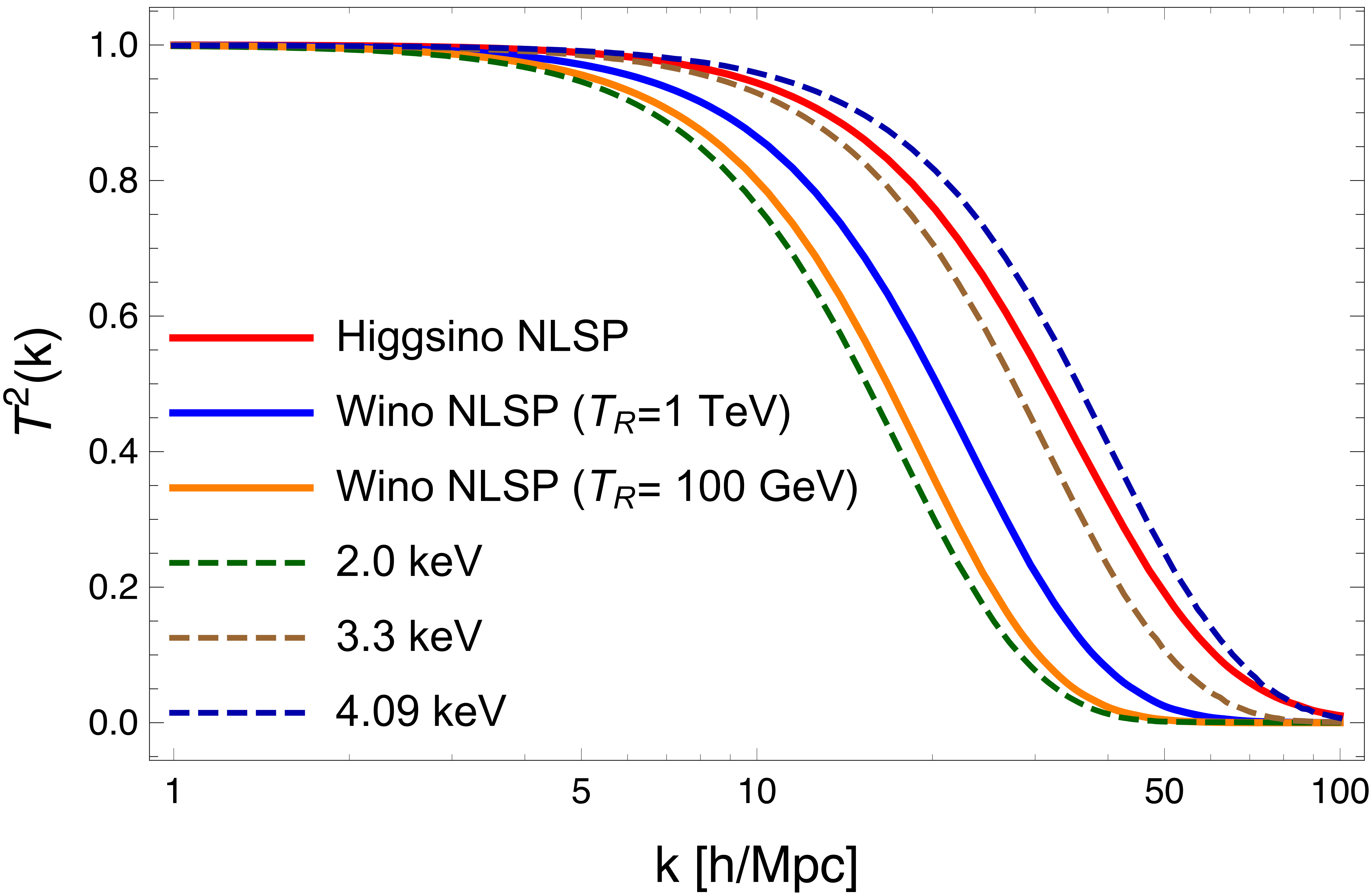
  ➡ need *linear matter power spectrum* to compare it with Ly-alpha

- Squared Transfer Function

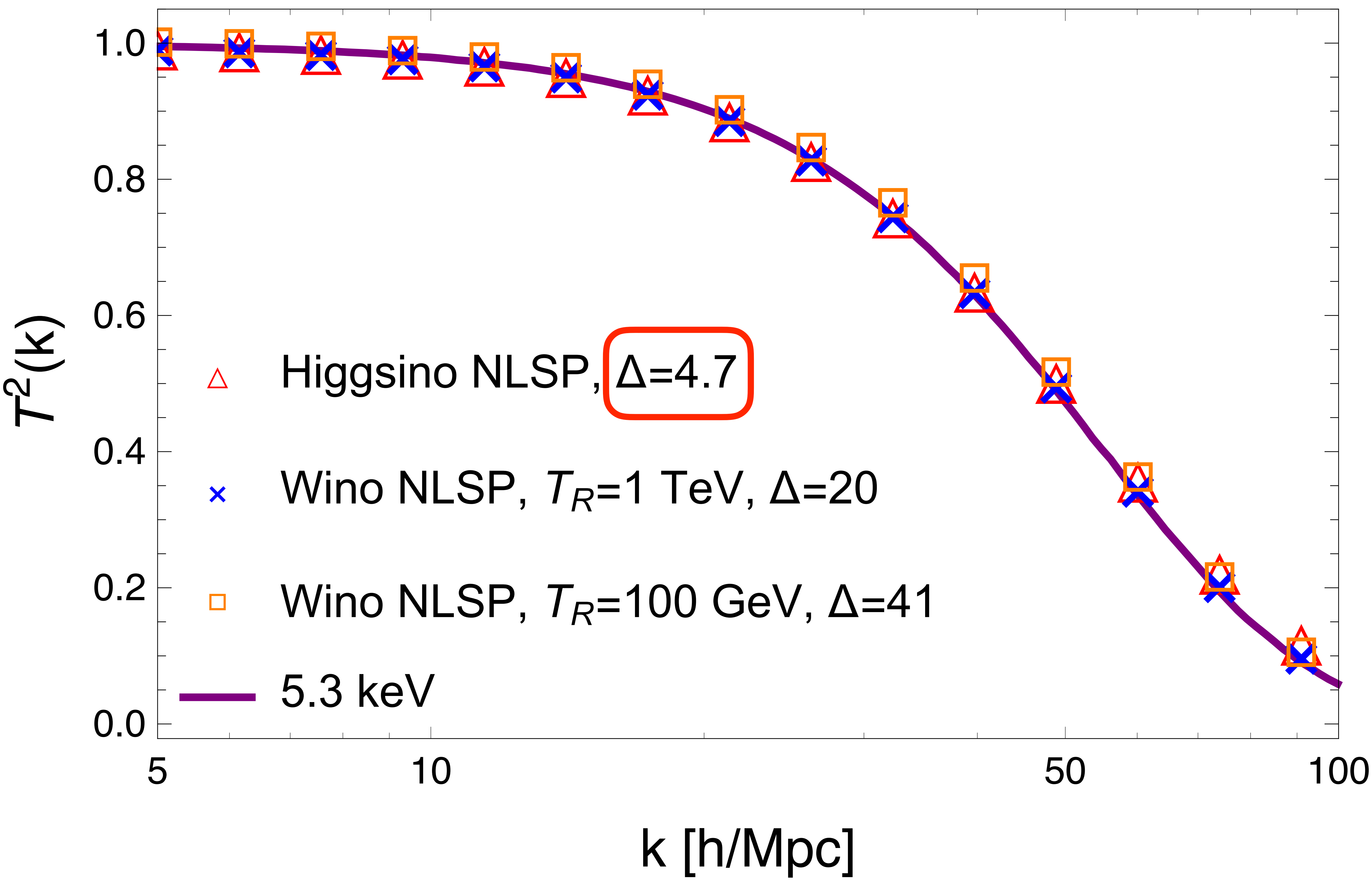
$$\mathcal{T}^2(k) = \frac{P(k)}{P_{\rm CDM}(k)}$$



MATTER POWER SPECTRUM



ENTROPY

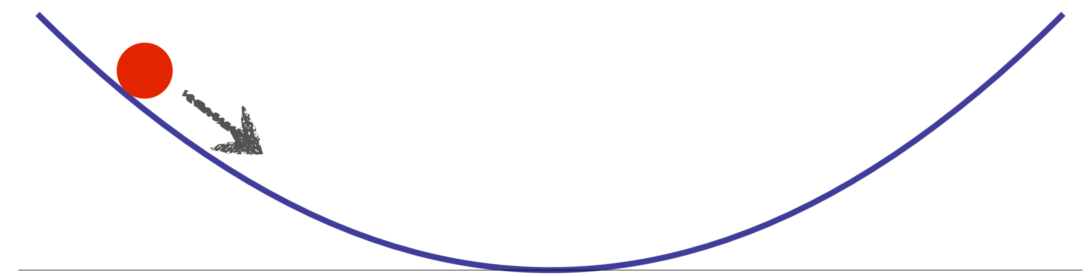


1.0
0.8
0.6
0.4
0.2
0.0
$T^2(k)$
5
10
50
100
k [h/Mpc]
Higgsino NLSP, Δ=4.7
Wino NLSP, T_R=1 TeV, Δ=20
Wino NLSP, T_R=100 GeV, Δ=41
5.3 keV

SAXION DECAY

- Coherent oscillation of saxion:

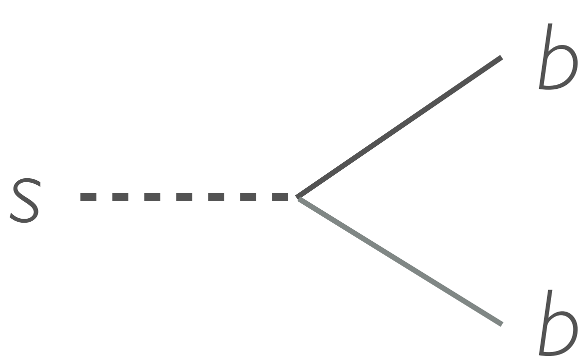
$$Y_s^{\rm CO} \simeq 1.9\times 10^{-6}\left(\frac{\rm GeV}{m_s}\right)\left(\frac{\min[T_R,T_s]}{10^7\,{\rm GeV}}\right)\left(\frac{s_0}{10^{12}\,{\rm GeV}}\right)^2$$



- Saxion dominated universe at

$$T_e^s = \frac{4}{3}m_s Y_s^{\rm CO} \simeq 2.5\times 10^2\,{\rm GeV}\left(\frac{\min[T_R,T_s]}{10^7\,{\rm GeV}}\right)\left(\frac{s_0}{10^{16}\,{\rm GeV}}\right)^2$$

- Saxion decay



$$v_{\rm PQ} = 2.5\times 10^{10}\ {\rm GeV}$$
$$m_s \simeq 110\ {\rm GeV}$$
$$T_D^s \simeq 53\ {\rm GeV}$$

$$\Rightarrow\quad \Delta = \frac{T_e^s}{T_D^s} \simeq 4.7$$

RELIC ABUNDANCE

- ***Higgsino NLSP (BM1) case***

By integrating $q^2 f_{\tilde{a}}(q)$

$$\Omega_{\tilde{a}} h^2 \simeq 0.1 \left(\frac{4.7}{\Delta}\right) \left(\frac{2.5 \times 10^{10}\,\mathrm{GeV}}{v_{\rm PQ}}\right) \left(\frac{m_{\tilde{a}}}{7\,\mathrm{keV}}\right)$$

Axinos can be the dominant DM.

- ***Wino NLSP case (BM2)***

hard to get enough entropy due to low T_R

For large T_R, it becomes similar to Higgsino NLSP.

CONCLUSIONS

- Axino can be DM when its mass is ~keV.
- Under R-parity violation, its decay signal can be observed: 3.5 keV X-ray excess?
- SUSY DFSZ model accommodates both mu-term for freeze-in production and bilinear RPV for axino decay.
- While 7 keV DM with FD dist. has tension with Ly-alpha m_{WDM}>3.3 keV, freeze-in axino can be consistent with m_{WDM}>5.3 keV with mild entropy dilution.
- Saxion (m=110 GeV) decay can make mild dilution.
- Axino can explain the observed DM density.